

Magnetic Field Configuration Models and Reconstruction Methods for Interplanetary Coronal Mass Ejections

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Abstract This study aims to provide a reference for different magnetic field models and reconstruction methods for interplanetary coronal mass ejections (ICMEs). To understand the differences in the outputs of these models and codes, we analyzed 59 events from the Co-ordinated Data Analysis Workshop (CDAW) list, using four different magnetic field models

Flux-Rope Structure of Coronal Mass Ejections

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and reconstruction techniques; force-free fitting, magnetostatic reconstruction using a numerical solution to the Grad–Shafranov equation, fitting to a self-similarly expanding cylindrical configuration and elliptical, non-force-free fitting. The resulting parameters of the reconstructions for the 59 events are compared statistically and in selected case studies. The ability of a method to fit or reconstruct an event is found to vary greatly; this depends on whether the event is a magnetic cloud or not. We find that the magnitude of the axial field is relatively consistent across models, but that the axis orientation of the ejecta is not. We also find that there are a few cases with different signs of the magnetic helicity for the same event when we leave the boundaries free to vary, which illustrates that this simplest of parameters is not necessarily always clearly constrained by fitting and reconstruction models. Finally, we examine three unique cases in depth to provide a comprehensive idea of the different aspects of how the fitting and reconstruction codes work.

Keywords Sun: corona · Sun: coronal mass ejections (CMEs)

1. Introduction

Coronal mass ejections (CMEs) are large-scale eruptions of plasma and magnetic flux with a typical size of 0.25 AU at 1 AU (*e.g.*, see Bothmer and Schwenn, 1998; Savani *et al.*, 2009; or Nieves-Chinchilla *et al.*, 2012). In the interplanetary space, they typically move at supersonic speeds between 200 and 2000 km s⁻¹, propagating from the solar corona to 1 AU in a few days. They can interact with the Earth’s magnetosphere, resulting in geomagnetic storms (Gold, 1959). Once CMEs drive through and interact with the interplanetary medium, they are often referred to as interplanetary CMEs (ICMEs). These ICMEs, measured *in situ*, for example at L1, may be composed of a fast forward shock, a dense sheath and ejecta material. Slow CMEs typically do not drive a shock. Hereafter, we use the term ICME once it is propagating in the interplanetary medium after leaving the corona (more or less after 10–20 R_{\odot}).

Interplanetary CMEs, measured *in situ*, are typically characterized by a low proton temperature, the presence of bi-directionally streaming electrons, unusual charge states of oxygen and iron, and various magnetic signatures (*e.g.* Lynch *et al.*, 2003; Zurbuchen and Richardson, 2006; Richardson and Cane, 2010). However, there is no single characteristic that is consistently observed. Lists of ICMEs measured by ACE and STEREO are maintained by Jian *et al.* (2006), Lepping *et al.* (2006), Richardson and Cane (2010), and Gopalswamy *et al.* (2008), among others. ICMEs have been classified into three different classes based on their magnetic field and plasma properties (Zurbuchen and Richardson, 2006):

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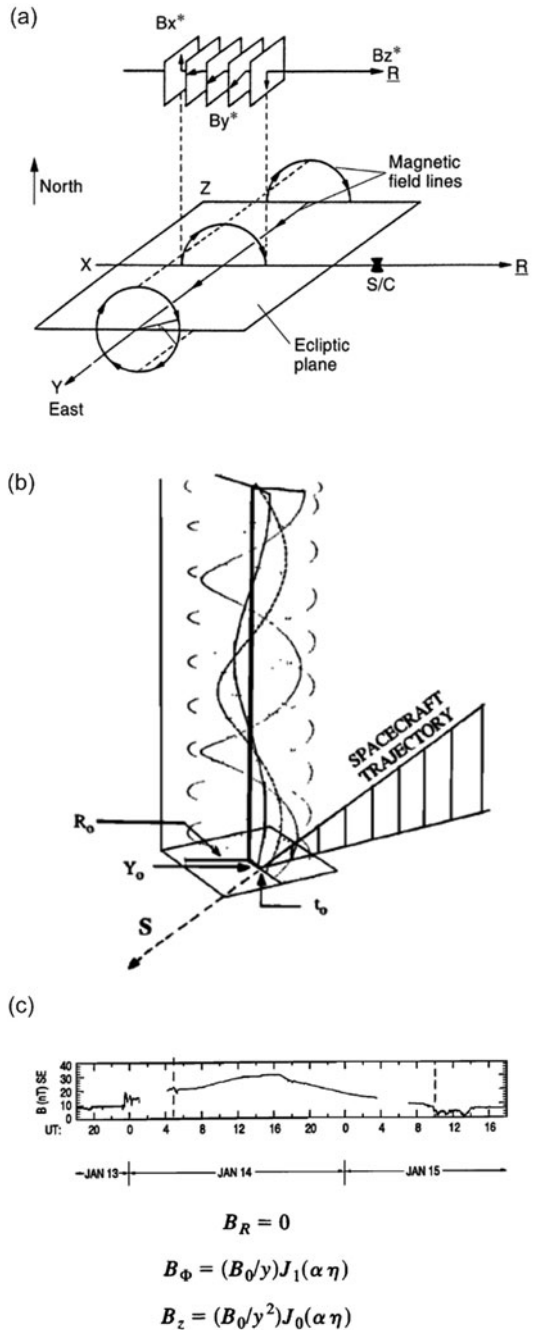
- i) Magnetic clouds (MCs): characterized by low plasma beta, low proton temperature and featuring a strong magnetic field with a smooth rotation (Burlaga *et al.*, 1981). To explain this last characteristic, Goldstein (1983) suggested that the magnetic field in an MC can be described by a force-free flux rope structure, where $\nabla \times \mathbf{B} = \alpha \mathbf{B}$. Burlaga (1988) assumed a constant α , in which case the flux rope satisfies the Lundquist solution (Lundquist, 1950) and can be expressed in terms of Bessel functions. This became the most widely used model to describe the magnetic field structure of MCs. MCs are believed to represent about one third of the total number of ICMEs (Gosling, 1990; Richardson and Cane, 2010). However, a number of other studies proposed that an even higher percentage of ICMEs are MCs; for example, the study by Li *et al.* (2011) and Marubashi (2000) estimated the proportion of MCs to be 50 % and 80 %, respectively. Other authors have studied the variation of the proportion of MCs among ICMEs during the solar cycle (Kilpua *et al.*, 2011; Richardson and Cane, 2010) and found that it varies from close to 100 % near solar minimum to about 20–25 % near solar maximum.
- ii) Ejecta or irregular and weak magnetic field ICMEs: These configurations have some but not all the properties of MCs. Typically, their magnetic field is too weak and/or irregular or has little rotation. Some of these ejecta can be additionally classified into subcategories, for example, “magnetic-cloud-like” (Zhang *et al.*, 2007) or magnetic flux rope (Rouillard, 2011). It has been proposed recently that some of these events are actually MCs whose legs passed through the observing spacecraft (Marubashi and Lepping, 2007), which has been confirmed with STEREO observations (Möstl *et al.*, 2010). According to these studies, ejecta are MCs but are not observed as such due to an observational bias.
- iii) Complex ICMEs (complex ejecta), which may result from the interaction of successive CMEs or from the interaction of CMEs with complex solar wind structures and streams (Burlaga, Plunkett, and St. Cyr, 2002; Wang, Ye, and Wang, 2003; Lugaz, Manchester, and Gombosi, 2005b). Double rotation within one MC has also been reported from *Ulysses* (Rees and Forsyth, 2004) and ACE observations (Steed *et al.*, 2011), showing that even isolated ICMEs may produce complex ejecta.

Understanding the nature of the magnetic field in ICMEs (magnetic clouds and ejecta) is a crucial part of comprehending their behavior and the way they affect the interplanetary medium and Earth’s magnetosphere throughout their propagation. Thus there have been several attempts to describe the structure of the magnetic field in ICMEs using models and magnetic field reconstruction codes. Figures 1 and 2 give some examples of proposed models.

To explain the behavior of the magnetic field in MCs, early models were developed to accommodate the definition of Burlaga *et al.* (1981) of an MC. Most of these models restrict the internal structure of the magnetic field to the force-free configuration where the magnetic field, \mathbf{B} , is described as $\nabla \times \mathbf{B} = \alpha \mathbf{B}$. Several magnetic field reconstructions and fitting models have been built upon this definition; among the most often used is the model by Burlaga (1988) and Lepping, Burlaga, and Jones (1990), which assumes a cylindrically symmetric solution and a constant α across the cloud.

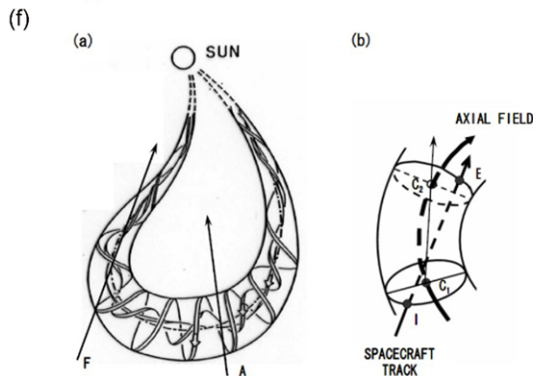
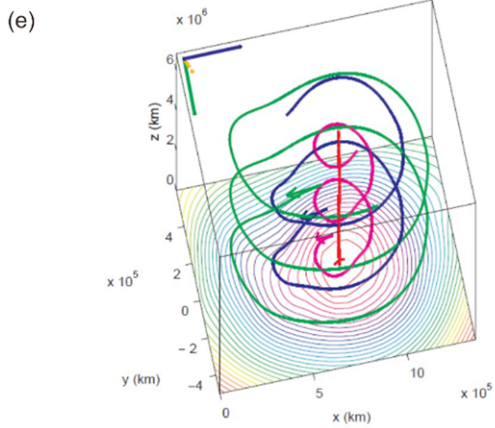
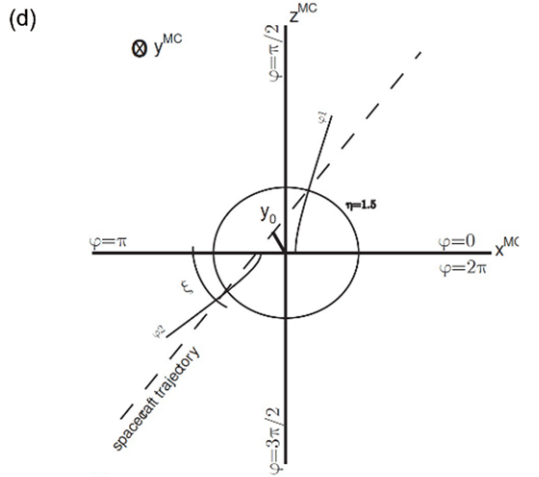
Marubashi (1986) also adopted a force-free model, but without assuming a constant α . Farrugia *et al.* (1993) and Farrugia, Osherovich, and Burlaga (1995) introduced a cylindrical model that incorporated self-similar expansion in two initially force-free constant- α codes, a cylindrical and a spherical model, and noted that the cylindrical configuration did not maintain the force-free state after it starts to expand. Thereafter, Shimazu and Vandas (2002) provided a modification of the mathematical formalism to keep the self-similarly expanding cylindrical model force-free.

Figure 1 Overview of different MC models. (a) Goldstein (1983) proposed that MCs are force-free structures, which was confirmed by Marubashi (1986) (figure adapted from Bothmer and Schwenn, 1998). (b) Lepping, Burlaga, and Jones (1990) succeeded in fitting the Lundquist model to *in situ* magnetic field observations after an idea by Burlaga (1988). (c) Farrugia *et al.* (1993) included expansion in the model.



Several other models have emerged thereafter that made different assumptions: fitting to a cylinder of elliptical cross-section by Hidalgo, Nieves-Chinchilla, and Cid (2002) and Hidalgo (2003); fitting to a kinematically distorted flux rope (Owens, Merkin, and Riley, 2006;

Figure 2 Overview of different MC models (*continued*). (d) Schematic representation of the spacecraft trajectory inside an MC of elliptical cross-section following the model of Hidalgo, Nieves-Chinchilla, and Cid (2002). (e) Hu and Sonnerup (2001) modeled magnetic flux ropes using a magneto-hydrostatic reconstruction technique without a predefined geometry. (f) Marubashi and Lepping (2007) included curvature in the classical model.



Vandas and Romashets, 2003; Vandas *et al.*, 2006; Démoulin and Dasso, 2009); non-cylindrical flux rope fitting (Mulligan and Russell, 2001; Owens *et al.*, 2012); torus fitting (Romashets and Vandas, 2003; Marubashi and Lepping, 2007). The Grad–Shafranov reconstruction technique (Hu and Sonnerup, 2002; Möstl *et al.*, 2009b) assumes a structure in magneto-hydrostatic (MHS) equilibrium with an invariant direction, and uses the Grad–Shafranov equation to describe the magnetic field in the structure. Therefore, this model does not only use the magnetic field measurements but also measurements of the plasma pressure, and it is only applicable to a structure possessing axial symmetry. The Grad–Shafranov technique has been tested using multi-spacecraft measurements (Liu *et al.*, 2008; Kilpua *et al.*, 2009; Möstl *et al.*, 2009a) and, recently, improvements to the algorithm have been presented by Isavnin, Kilpua, and Koskinen (2011).

While most techniques reconstruct or fit ICMEs as some type of twisted flux ropes, this result is not necessarily well tested. It has recently been shown by Al-Haddad *et al.* (2011) that the Grad–Shafranov technique is designed in a way that it will always reconstruct a helical flux rope from a rotating magnetic field observed by a single spacecraft, even if the magnetic field is not a helical flux rope in 3D. This is precisely the reason why the Grad–Shafranov reconstruction assumes an invariance along the cloud axis (so-called 2.5D method). As was shown in Al-Haddad *et al.* (2011), this is equivalent to assuming a helical solution from the beginning. The assumption of invariance along the cloud axis is also made in all the other fitting techniques used in this article. Other recent observations pointing toward a more complicated geometry of ICMEs at 1 AU have been made by Kahler, Krucker, and Szabo (2011), Farrugia *et al.* (2011) and Nieves-Chinchilla *et al.* (2012).

Using 2D and 3D numerical simulations, it is possible to show that a CME initiated at the solar surface (for example using a twisted flux rope) evolves through its interaction with the solar wind into a typical ICME (Riley, Gosling, and Pizzo, 1997; Riley *et al.*, 2003; Odstrčil and Pizzo, 1999; Manchester *et al.*, 2004; Chané *et al.*, 2006; Shen *et al.*, 2007). This idea has now been confirmed by STEREO observations (see, for example, Savani *et al.*, 2010). In these simulations, synthetic satellite measurements made at 1 AU show a typical ICME structure with a fast shock, preceding a dense sheath and an ejecta. Synthetic coronagraphic and heliospheric images are also able to reproduce typical views of CMEs (Lugaz, Manchester, and Gombosi, 2005a; Manchester *et al.*, 2008; Riley *et al.*, 2008; Lugaz *et al.*, 2009; Odstrčil and Pizzo, 2009).

Riley *et al.* (2004) compared magnetic field reconstructions and fitting models for ICMEs by fitting two different time series of a simulated ICME with five different techniques: three force-free models, the elliptical model by Hidalgo, Nieves-Chinchilla, and Cid (2002), and the Grad–Shafranov reconstruction code. The results of the reconstructions were then compared again to the 2.5D MHD simulation. One of the difficulties that the authors encountered was choosing the boundaries. Another difficulty was recognizing an MC within the provided time series. It is particularly hard to observe the properties characterizing an MC when the minimum variance technique is used and the impact parameter is large.

Three of the models used in the study by Riley *et al.* (2004) were also used in the present study. We also added the self-similarly expanding cylindrical fitting of Marubashi and Lepping (2007) and only kept one classical – no expansion – force-free model. However, it is important to stress that in this project we applied the comparison to a list of observed events. Here, the authors address some of the problems discussed in Riley *et al.* (2004) by performing more detailed analyses and proposing fixed boundaries.

In this study, we compare the reconstruction of four different magnetic field fitting techniques for most of the fast ICMEs and MCs observed during Solar Cycle 23, for which a source region can be identified (see details in Section 2). Our goals are to acquire a broader

understanding of the magnetic field structure in MCs and to investigate if there are statistical differences between the codes for MCs and non-MC ejecta, and take steps toward deciphering if all ICMEs can be expressed as flux ropes if a sufficiently sophisticated model is used. We do this by comparing the results of the different fitting techniques. In Section 2, we describe the data used in this study followed by a succinct description of the different codes. In Section 3, we discuss the full results for the 59 selected events and in Section 4, we present two events for which we used the same boundaries for all codes. We conclude and discuss our results for the nature of magnetic fields in ICMEs in Section 5.

2. Data and Model Overview

2.1. CDAW Data

The data are taken from the list of the Coordinated Data-analysis Workshop (CDAW) of MCs and ejecta during Solar Cycle 23. The list can be found at http://cdaw.gsfc.nasa.gov/meetings/2010_fluxrope/LWS_CDAW2010_ICMEtbl.html.

This list contains all shock-driving ICMEs from Solar Cycle 23 for which it was possible to determine the source region, and for which the source region was within $\pm 15^\circ$ of the disk center. This list is a subset of the list of all interplanetary shocks observed during Solar Cycle 23 from Gopalswamy *et al.* (2010). The list contains 59 ICMEs classified as follows: 24 events classified as magnetic clouds (MCs) and 35 events classified as ejecta (EJ).

2.2. Grad–Shafranov Reconstruction Technique

For the Grad–Shafranov (GS) reconstruction, we used the code by Hu and Sonnerup (2002) and followed the guidelines for using it discussed by Möstl *et al.* (2009b). This is based on magneto-hydrostatic equilibrium of a system with an invariant direction and is a solution for what is basically a numerical boundary problem. The GS equation is

$$\frac{\partial^2 A}{\partial x^2} + \frac{\partial^2 A}{\partial y^2} = -\mu_0 \frac{dP_t(A)}{dA}, \quad (1)$$

where $A(x, y)$ is the magnetic vector potential and P_t is the transverse pressure defined as $P_t = p + B_z^2/(2\mu_0)$, with p being the thermal pressure. The total magnetic field is given by $\mathbf{B} = \nabla A \times \mathbf{e}_z + B_z \mathbf{e}_z$. The first step to solving this equation is to determine the invariant axis, z . This is done first using a minimum variance analysis on the magnetic field components and by finding a frame in which the transverse pressure is a single-value function of the magnetic vector potential, A . After this, the GS equation can be solved numerically.

The main assumptions of the technique are that the structure is assumed to be invariant along the ICME axis (so-called 2.5D), and that it is assumed to be time-independent during the whole measurement by the spacecraft. Compared to other techniques, it does not include expansion or a toroidal geometry. One of the potential problems with the Grad–Shafranov technique is the stability of the integration and the influence of the solver and stabilization procedure on the results. However, the Grad–Shafranov reconstruction has been tested by using observations of ICMEs by two spacecraft at a significant separation with respect to the ICME size (Liu *et al.*, 2008; Möstl *et al.*, 2009a, 2009b). Some of the stability problems have also recently been addressed in Isavnin, Kilpua, and Koskinen (2011).

One of the main advantages of the GS method is that the shape of the flux rope's cross-section and the number of flux ropes inside the reconstructed interval are both not pre-defined, but are an output of the technique. Another advantage is that the boundaries of the ICMEs are chosen through an optimization procedure to render the $P_t(A)$ function as close to single-valued as possible. The $P_t(A)$ function is fitted with a polynomial function with exponential tail(s), and only the events for which the fitting residue is low (typically less than 0.1) and for which the fitting appears to be visually correct are deemed successfully reconstructed. More details can be found in Hu *et al.* (2004).

The GS reconstruction gives the following results: flux rope orientation (longitude and latitude), impact parameter (closest approach of the spacecraft to the MC axis), magnetic flux (axial and poloidal), axial current, the cross-section shape and the central field strength. The GS reconstruction feasibility was checked for every event in the CDAW list. We primarily used data from the *Wind* spacecraft (MFI (Lepping *et al.*, 1995) and SWE (Ogilvie *et al.*, 1995) instruments), and switched to ACE (MAG (Smith *et al.*, 1998) and SWEPAM (McComas *et al.*, 1998) instruments) if there was a data gap at *Wind*. The GS reconstruction method was successful for

- full list: 20 events out of 59 (34 %),
- 17 MC events out of 24 (71 %),
- 3 EJ events out of 35 (8 %).

This means that the applicability of using GS reconstruction is closely tied to the definition of a magnetic cloud (higher-than-average field strength, smoothly rotating magnetic field vector, low β , low T_p). Typically, this means that for these cases, one magnetic field component (B_y or B_z) needs to be bipolar and the other unipolar, depending on the orientation of the flux rope (see discussion in Bothmer and Schwenn, 1998). Moreover, we were able to reconstruct about 1/3 of the events in the complete list, which is similar to the often-quoted ratio between all ICMEs and those that contain parts that satisfy the MC definition.

2.3. Self-similarly Expanding Cylinder Fitting Technique

This model is a force-free model that assumes a spatially constant α and performs the fitting in cylindrical coordinates. Here, α is chosen fixed for every event and equal to 2.4. The model takes into account self-similar expansion, as was initially suggested by Farrugia *et al.* (1992, 1993). The original inclusion of the self-similar expansion has been found to cause a deviation from the force-free state during expansion. Some refinements proposed by Shimazu and Vandas (2002) have been incorporated into the model to preserve the force-free state for the duration of the MC travel. As a result, the magnetic field intensity B and the radius r of the fitted MC at a given time as produced by the model are defined as follows:

$$\mathbf{B} = \frac{B_0}{\left(1 + \frac{t}{t_{\text{exp}}}\right)^2}, \quad (2)$$

$$\mathbf{r} = r_0 \left(1 + \frac{t}{t_{\text{exp}}}\right), \quad (3)$$

where B_0 is the magnetic field intensity on the axis of the CME cylinder, t_{exp} is the expansion time, t is the time since the first encounter of the spacecraft with the CME, and r_0 is the distance between the cylinder axis and the spacecraft at the time of the first encounter.

The model also provides values for the longitude and latitude of the axis, as well as the impact parameter, which is defined as the distance between the closest axis of the cylinder and the path of the spacecraft. The model simultaneously fits the following quantities for both chiralities (positive and negative): B_0 , r_0 , t_{exp} , the longitude, latitude, and impact parameter. The final selection was made by choosing the sign of the chirality that gives the best fit. The details of the model can be found in Appendix A of Marubashi and Lepping (2007). The self-similarly expanding cylinder fitting method was successful for

- full list: 50 events out of 59 (85 %),
- 22 MC events out of 24 (91.5 %),
- 28 EJ events out of 35 (80 %).

There were seven events that were found to be contracting (negative t_{exp}) and three events whose expansion time was longer than 500 hours, which corresponds to less than 10 % change in the axial magnetic field strength over the crossing of a typical magnetic cloud (~ 1 day).

2.4. Force-Free Constant- α Reconstruction

This force-free fitting model is based on the force-free flux constant- α (FFCA) rope developed by Burlaga (1988) and subsequently optimized by Lepping, Burlaga, and Jones (1990). In our model, α is fixed to equal 2.4. This implies that the flux rope's magnetic field on the outer shell is almost completely poloidal. The value of 2.4 is not optimized and could be changed (*i.e.*, reduced in value) if it was believed that on the outer shell of the flux rope the field lines were still helical. This might be the case if the rear or front edge of the flux has been “eroded” away from its ideal case (perhaps by reconnection), as proposed by Dasso *et al.* (2006) and recently confirmed by Ruffenach *et al.* (2012). The α parameter has also been used as a free variable in some models (*e.g.* Lepping, Burlaga, and Jones, 1990). Since α is constant and fixed to equal 2.4 for the FFCA and for the self-similarly expanding cylinder models, it is clear that these two models are identical when the expansion is equal to 0.

The flux rope axis was calculated by a minimum variance analysis (MVA). This is a technique originally developed for solving the normals to a tangential discontinuity (TD) (Sonnerup and Scheible, 1998). To obtain the orientation of an ICME, the axis is determined by the intermediate eigenvector, as pointed out by Goldstein (1983).

The chirality of the flux rope is determined by the sense of direction in which the magnetic field rotates in the MVA reference frame. The impact parameter and magnetic field magnitude are the only two free variables. They are optimized using a nonlinear optimization routine (Nelder and Mead, 1965). For each event, the fitting was attempted with two different boundaries: those from Lepping *et al.* (2006) and those selected by K. Marubashi. In addition, for eight events, a different set of boundaries was chosen directly by N.P. Savani because the other two sets of boundaries were judged to be unsatisfactory for these events. Overall, the method was successful for

- full list: 55 events out of 59 (93 %),
- 24 MC events out of 24 (100 %),
- 31 EJ events out of 35 (88.5 %).

2.5. Elliptical Cross-Section Fitting Model

The elliptical, non-force-free model (Hidalgo, Nieves-Chinchilla, and Cid, 2002; Hidalgo, 2003) considers that ICMEs have elliptical cross-section due to their interaction with the

solar wind. In this model, the expansion of the local cross-section is taken into account. Two characteristics distinguish this model from the others used in this paper: i) non-force-free condition is assumed, and, ii) an elliptical cylindrical coordinate system is chosen to solve the Maxwell equations. The model assumes a flux rope magnetic field geometry and treats the radial and axial electric currents separately. The local solutions of the Maxwell equations provide the three magnetic field components that are to be transformed into spacecraft coordinates. The only assumptions are that the radial and axial components of the current density are constant. The cross-section of the flux ropes is not required to be circular. This model was compared with the circular cross-section model of Lepping, Burlaga, and Jones (1990) in Nieves-Chinchilla, Hidalgo and Sequeiros (2005) and compared with information from remote-sensing measurements in Nieves-Chinchilla *et al.* (2012).

The ellipticity, the current density components, the expansion time of the cloud, and the axial magnetic field are determined by fitting the magnetic field data as explained in Section 2 of Hidalgo (2003). Multiple regression analysis is used to infer the spacecraft trajectory through the flux rope. The method converged for all events. However, the fit goodness was reported with a flag: 0 = poor; 1 = good; and, 2 = very good. This “goodness” is based on the correlation coefficient, fitting residue, and visual goodness. In the rest of the analysis, we take into account only fitting results with quality 1 or 2 (*i.e.*, removing poor fits). In that case, the elliptical cross-section fitting method was successful for

- full list: 39 events out of 59 (66 %),
- 24 MC events out of 24 (100 %),
- 15 EJ+ events out of 35 (43 %).

2.6. Selection of the ICME Boundaries

Selecting the boundaries of an ICME is still a challenging problem. This is because there is no single definitive characteristic of ICMEs. Some authors based the ICME boundaries purely on magnetic properties; others used a low proton temperature as one of the criteria; while others used charge state information. Many of the *in situ* signatures have been taken into account to determine the boundaries of ICMEs. However, each of these indicators provided different boundaries, which leaves the problem still unsolved. In this study, we let each group determine the boundaries for the ICMEs. For three events, we re-ran the analysis using pre-set boundaries as given by the self-similarly expanding cylinder fitting technique.

3. Comparative Analysis of the Fitting Parameters – Results

Out of 59 events on the CDAW list, the magnetic field was successfully reconstructed for 20 events by the GS technique, mostly for MCs. By comparison, the FFCA fitting and the self-similarly expanding cylinder methods were able to fit more than 85 % of the list, and the elliptical cylinder fitting method 66 %. In the following sections, we compare the fitted parameters from different codes with those from the FFCA code because i) it is the simplest code in our study, and ii) it succeeded in fitting most of the cases with the least deviation from the results of the other models. The number of cases that were analyzed in common by the FFCA code and the other codes is as follows: 20 with the GS reconstruction technique, 45 with the self-similarly expanding cylinder fitting method, and 37 with the elliptical cylinder fitting method.

3.1. Chirality

For all the 20 events reconstructed by the GS code, the chirality (sign of the magnetic helicity) is the same as for the self-similarly expanding cylinder and the FFCA model. There are three events for which the FFCA and self-similarly expanding cylinder found opposite signs for the helicity (events 13, 15, and 52 in the CDAW online list). This shows that even the simplest result of fitting and reconstruction methods is not necessarily well constrained. The difference is primarily due to the choice of boundaries. In two of these three cases (13 and 15), when the FFCA fit was performed again with a different choice for the boundaries, it returned the same sign of helicity as that found by the self-similarly expanding cylinder. In fact, it is worth noting that the chirality found with the FFCA differs for seven events depending on whether the boundaries are chosen following Lepping *et al.* (2006) or chosen specifically for this study by K. Marubashi.

One event is more ambiguous. For event 52, the FFCA fit is relatively poor and gives left-handed ejecta for the boundaries selected by K. Marubashi as well as for those from Lepping *et al.* (2006), whereas the self-similarly expanding cylinder fit gives right-handed ejecta. We note that the best-fit axial magnetic field by FFCA for this event is less than 6 nT for the two sets of boundaries, reflecting very weak ejecta.

This overall result about the chirality may appear inconsistent with what was found in previous studies by Dasso *et al.* (2003, 2006), who found the magnetic flux and helicity to be nearly independent of the fitting and reconstruction model for two well-observed MCs. Here, we find that this is generally true and consistent with our results. However, we also show that for ejecta whose boundaries are not clearly defined because of incomplete rotation, fittings with different boundaries can yield opposite chirality for the same event.

3.2. Maximum Magnetic Field Strength

For all cases in common between the FFCA code and one of the other codes in this study, we calculated the correlation of the fitted axial magnetic fields. For the self-similarly expanding cylinder fitting method, we used the axial magnetic field halfway through the cloud, corresponding approximately to the magnetic field at the closest approach to the cylinder axis, as calculated from Equation (2). We did not compare the value of the axial magnetic field for the elliptical cross-section fitting model with that of the other models. This is because the axial magnetic field in this model is the sum of two fields: the magnetic field created by the currents and the axial field that is a parameter of the model. Figure 3 summarizes the magnetic field results.

Overall, the correlation between the FFCA, GS, and self-similarly expanding cylinder code is quite good (~ 0.84). The axial magnetic field value from the GS code tends to be slightly higher than that from the FFCA especially for stronger magnetic fields, whereas the self-similarly expanding cylinder code tends to return higher values for the magnetic field for weak magnetic fields.

3.3. Orientation of the ICME Axis: Longitude and Latitude

The orientation of the ICME as it arrives at Earth is an essential parameter because it is often used to determine rotation of the CME as it propagates in the heliosphere (Yurchyshyn *et al.*, 2001). In addition, precisely knowing the orientation of an ICME is particularly important because the orientation of an ICME is known to be related to its geo-effectiveness (see, for example Zhao and Hoeksema, 1998). As can be seen in Figure 4, the orientation of some

Figure 3 Axial magnetic field for the Grad–Shafranov (red diamonds) and self-similarly expanding cylinder (purple stars) models compared with the axial magnetic field for the FFCA model. The red and purple dashed lines show the best linear fits to the Grad–Shafranov and self-similarly expanding cylinder models.

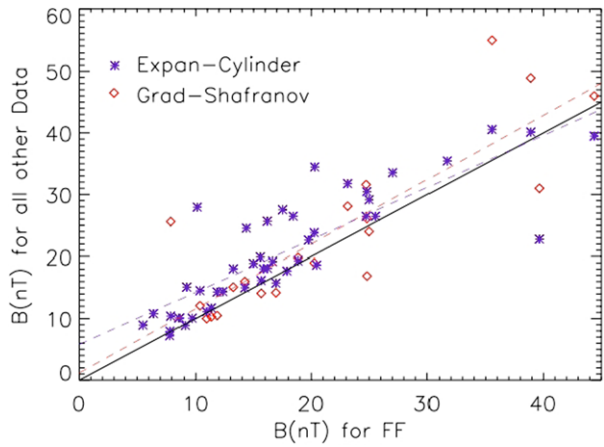
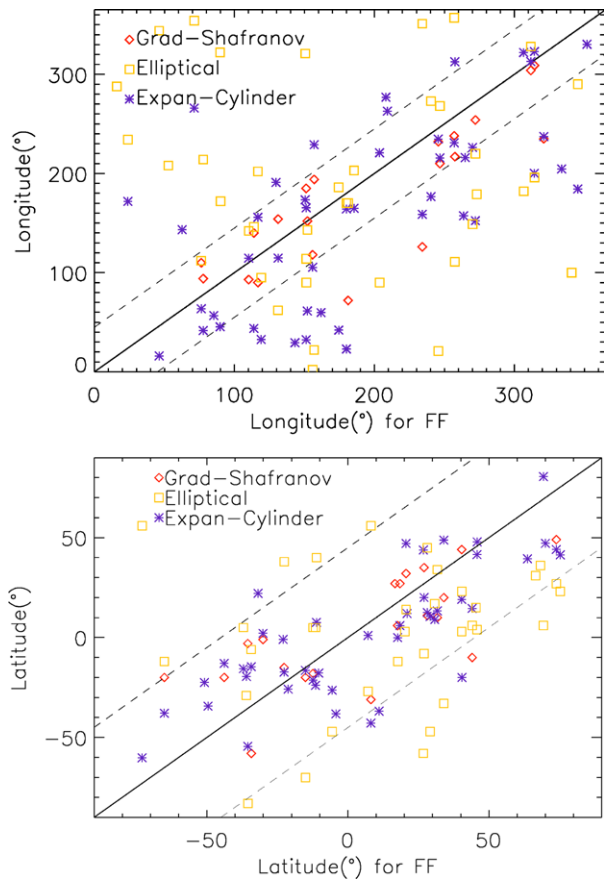


Figure 4 Same as Figure 3 but for the longitude and latitude of the reconstructed ICME axis. The yellow squares are for the elliptical cross-section fitting model. The dashed lines highlight the angles $\pm 45^\circ$ from the FFCA angle.



ICME events differs greatly depending on which model is used. Indeed, only for one CME (event 44) did all events give an orientation of the axis within $\pm 45^\circ$ of each other.

First, we focus on a comparison of the orientation between the Grad–Shafranov, self-similarly expanding cylinder and FFCA models for the 20 events successfully reproduced by all three methods. For seven events, the three methods give an orientation within $\pm 45^\circ$ of each other. For eleven other events, the Grad–Shafranov reconstruction technique gives an orientation of the axis within $\pm 45^\circ$ of one of the other two methods (but not the other). In fact, although the FFCA and self-similarly expanding cylinder fitting methods are closely related, the axis of the fitted cylinder is found to agree for only about 45 % of the events (21 out of the 45 common cases; for 14 of these, the orientation is within $\pm 30^\circ$). While this may appear surprising considering the similarities between the two methods, it should be noted that the force-free method without expansion relies on the minimum-variance analysis to obtain the orientation of the flux rope, whereas the expanding cylinder method relies on a multi-parameter fitting procedure to minimize the deviation between the model and the measurements. In addition, for many events, the two methods used different boundaries.

Comparing the direction obtained from the elliptical cross-section model with that from the force-free cylindrical models (with and without expansion), we find that the direction from all three models only agree for one event. For seven additional events, the direction from the elliptical cross-section model is within $\pm 45^\circ$ of the direction from the FFCA fitting model; for three additional cases, the same is true as compared to the direction from the self-similarly expanding cylinder model, and for one event it is true as compared with the direction for the Grad–Shafranov reconstruction technique.

Overall, there are only 34 events reconstructed by three or more methods. For 30 of these events, the orientation of the axis of the reconstructed ICME is consistent between at least two methods. The same is only true for three out of the 18 events reconstructed by two methods only. Finally, seven events were reconstructed by one method only.

3.4. Difference Between MCs and Non-MC Ejecta

Here, we briefly compare the average values of some parameters for the MCs with those for the non-MC ejecta. We only focus on the FFCA and self-similarly expanding cylinder, because the Grad–Shafranov reconstruction method is only able to reconstruct three non-MC ejecta (out of 35) and the elliptical cross-section fitting has a different set of parameters than these two models. Consistent with their definitions, non-MC ejecta are found to have, on average, a weaker magnetic field than MCs: 13.75 nT vs. 22.1 nT (FFCA) and 17.7 nT vs. 34.4 nT (expanding cylinder). We also find that the average impact parameter is larger for non-MC ejecta than for MCs: 0.73 vs. 0.54 (expanding cylinder) and 0.32 vs. 0.29 (FFCA). This may indicate that ejecta are indeed flux-rope-like magnetic clouds but are measured at a large impact parameter, resulting in a weaker magnetic field and less clearly defined properties. The average impact parameter for the events reconstructed by the GS method is 0.29 (FFCA) and 0.47 (expanding cylinder), once again showing that the GS method was only successful for MCs or MC-like ejecta.

4. Detailed Analyses of Three Events

We focus on events 10, 20, and 27, which we fitted using the same boundaries. These were the events chosen for the self-similarly expanding cylinder fitting method.

4.1. Event 10

This event is a left-handed non-MC ejectum observed in 26–28 June 1999. This was one of the events resulting in a poor fit for the elliptical cross-section model and was not initially

Table 1 Intervals used by different researchers for the events 10, 20, and 27. In bold are given the intervals chosen for events compared with the same boundaries. In parentheses are given the day of the month corresponding to the interval boundary. Events 10, 20, and 27 took place in June 1999, July 2000, and November 2000, respectively.

Event	10 (June)	20 (July)	27 (Nov.)
GS	N/A	08:36(27)–19:24(27)	23:56(06)–17:20(07)
FFCA	06:00(26)–22:00(26)	07:30–19:30	22:10(06)–17:10(07)
Exp. Cylind.	06:00(26)–19:00(26)	07:30–19:30	22:10–17:10
Ellip	22:36(26)–02:36(27)	10:10–20:40	23:08–17:45

successfully reconstructed by the Grad–Shafranov method. For this method, one problem was the presence of a shock propagating inside the ICME at 19:16 UT on 26 June. Because the GS reconstruction method attempts to find the best orientation of the ICME by requiring the function $P_i(A)$ to be single-valued, the presence of a shock in the magnetic and plasma pressure in one side of the cloud makes this an impossible task. This has been previously noted in Kilpua *et al.* (2009) and Möstl *et al.* (2009a). It is not a problem encountered by the other fitting methods.

Initially, some researchers found two possible flux rope intervals: one shortly after the shock wave on 26 June lasting about 16 hours (referred to hereafter as 10a) and one corresponding to the initial interval fitted with the self-similarly expanding cylindrical code of 6.5 hours at the end of 27 June and the beginning of 28 June (referred to hereafter as 10b from 20:30 UT on 27 June to 03:00 UT on 28 June). We discuss here the results from the longer, first interval. The intervals found by different researchers are shown in Table 1.

The analysis was performed again for interval 10a using the interval of 26 June from 06:00 to 19:00. In Figure 5, we show the fitting results for the different methods corresponding to this interval. Using the same interval for all the methods, the magnitude of the magnetic field in the ICME is found to agree relatively well between the two force-free and the Grad–Shafranov methods; it corresponds to a value $\sim 15–20$ nT. The orientation of the cloud is found to be consistent with a latitude of $0^\circ \pm 12^\circ$ and a longitude of $65^\circ \pm 25^\circ$ between these three methods, whereas the elliptical cross-section model finds a value about 90° away. This example illustrates the importance of the choice of boundaries for the study of magnetic clouds and ICMEs, because three of the four codes find very similar parameters for same boundaries.

4.2. Event 20

This event is a left-handed non-cloud ejectum observed on 27 July 2000. It is one of the three nonmagnetic cloud ejecta successfully reconstructed with the Grad–Shafranov reconstruction method. Initially, as we let each researcher select the ejectum boundaries, the direction of the cloud axis agreed well between the four methods: a longitude of $212^\circ \pm 32^\circ$ and a latitude of $22^\circ \pm 12^\circ$. All methods were within 45° of each other, except for the self-similarly expanding cylinder and the elliptical cross-section cylinder, which were about 55° apart.

The impact parameter was found to be close to 0 for the Grad–Shafranov method, but larger than 0.5 for the other three methods. The magnitude of the axial magnetic field was about 20 nT for the Grad–Shafranov method and about 8 nT for the two force-free constant-alpha models.

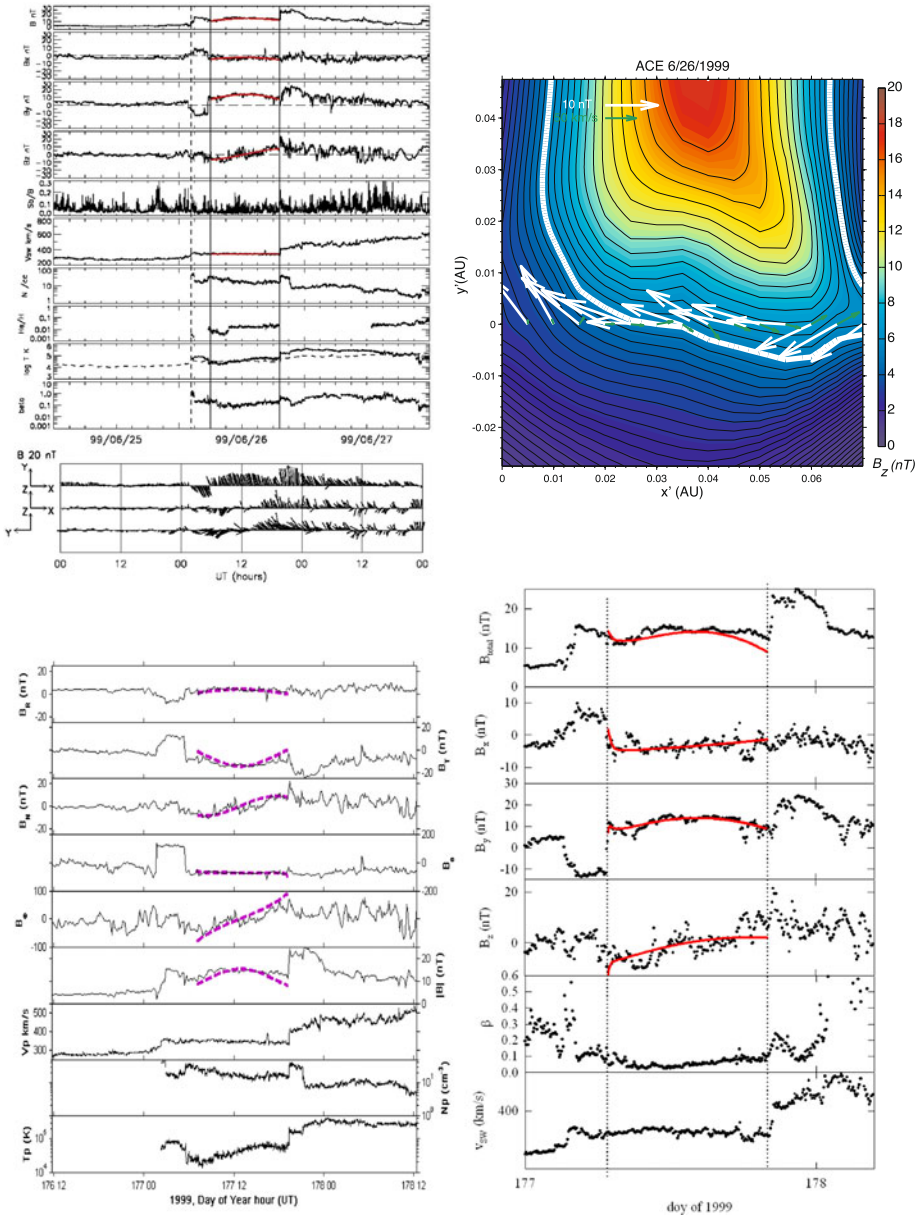


Figure 5 Fitting for event 10a by the different codes. Top left: Self-similarly expanding cylinder code fit shown in red. The different panels are the total magnetic field, B_x , B_y , B_z , the magnetic field fluctuation, the solar wind speed, proton density, alpha-to-proton ratio, and plasma β from top to bottom. Bottom left: FFCA code fit shown in purple. The panels show from top to bottom the radial, tangential, normal (R, T and N) components, the longitude and latitude angles (in RTN) and the magnitude of the magnetic field, the proton velocity, density, and temperature. Top right: GS reconstruction map, the axial magnetic field is color-coded and the black contours are magnetic field lines in the plane of the cross-section of the ICME. Observed magnetic field components in this plane are shown with white vectors and velocity is depicted in green. Bottom right: Elliptical cross-section fit in red. The different panels show the magnetic field, its three components, the plasma β , and the solar wind speed from top to bottom.

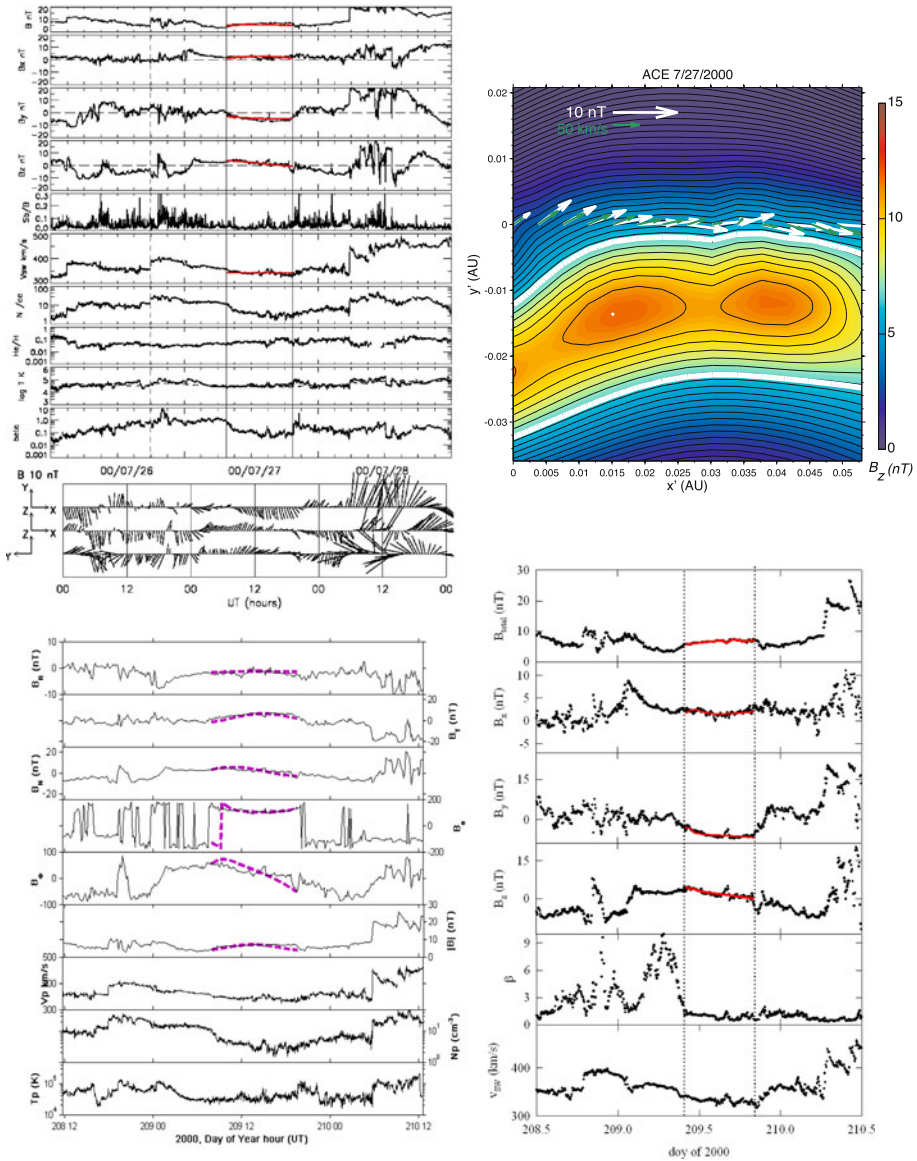


Figure 6 Same as Figure 5 but for event 20.

As can be seen from Table 1, the different researchers chose boundaries very close to each other, and choosing exactly the same boundaries does not drastically change the results, except that the maximum axial magnetic field for the Grad–Shafranov model was 12 nT with an impact parameter of about 0.25, more consistent with the values found by the other models. In Figure 6, we show the results of the fit using the boundaries chosen for the self-similarly expanding cylinder fit.

4.3. Event 27

This event is a left-handed magnetic cloud observed on 6–7 November 2000. All methods agree on the maximum magnetic field strength (~ 25 nT for Grad–Shafranov and the FFCA model and ~ 31.5 nT for the self-similarly expanding cylinder). However, there is initially little agreement regarding the orientation of the magnetic cloud with a latitude of $14^\circ \pm 14^\circ$ and a longitude of $145^\circ \pm 55^\circ$ for the four methods. As for the previous event, the boundaries chosen independently by different researchers were relatively consistent. One significant difference is the inclusion or exclusion of a region of strong magnetic field without much rotation at the “back” of the MC (until 02:40 on 8 November). Previous studies (Dasso *et al.*, 2006, 2007; Möstl *et al.*, 2008) have discussed how reconnection between an MC and the ambient solar wind during the ICME propagation may result in this type of region which strictly speaking belongs to the MC. Each researcher decided exactly where to end the MC. Since there is no counterpart flux on the MC front to the flux on the back of the MC, it is probably better not to take this region into account.

Using the same boundaries, there is a better agreement that the magnetic cloud has a low inclination (latitude $\sim 5^\circ$) but an only slightly better agreement about the precise orientation or the value of the impact parameter. The fits and reconstructions are shown in Figure 7.

5. Summary and Conclusions

In this article, we have presented a comparison of the reconstruction results of 59 ICMEs measured *in situ* during Solar Cycle 23 using four different reconstruction or fitting methods. The events our study focused on were chosen because their source region was known and was within $\pm 15^\circ$ of the central meridian. Our data included 24 events identified as magnetic clouds (MCs) following the definition of Burlaga *et al.* (1981) and 35 non-MC ejecta. We find that the two force-free techniques (with or without expansion) are able to fit the vast majority of the events ($\geq 85\%$) including more than 75% of the ejecta. The Grad–Shafranov reconstruction method, which assumes magneto-hydrostatic equilibrium and uses the plasma pressure is found to only work for clear magnetic clouds, because it was only successful in reconstructing 8% of the non-cloud ejecta (3/35). In addition, the Grad–Shafranov reconstruction is not able to reconstruct ICMEs into which a shock is propagating because it violates its initial assumption. However, one of the advantages of this method is the fact that the boundaries of the ICME are a result of the reconstruction, and they do not need to be chosen before performing the fits, as for the other methods. The expanding elliptical cross-section was able to fit about two-thirds of the events with a reasonably good result (all magnetic clouds and about 40% of the ejecta).

We have found that even the ICME chirality (sign of the magnetic helicity), one of the simplest possible reconstruction parameters, is not necessarily clearly constrained, and opposite chirality can be found for the same event because of differences in the choice of the boundaries. In most cases, the choice of boundaries alone explains the difference between different methods.

Regarding the maximum magnetic field strength, we have found that, typically, the results of the Grad–Shafranov method, the FFCA, and self-similarly expanding fittings are well correlated and agree to within about 25%. The Grad–Shafranov method returns higher values than the FFCA fitting method for strong magnetic fields, whereas the force-free model, which includes self-similar expansion, returns higher values than the FFCA fitting method for weak magnetic fields. Gulisano *et al.* (2005) previously compared the maximum magnetic field at the cloud axis for 20 magnetic clouds using four fitting methods and found

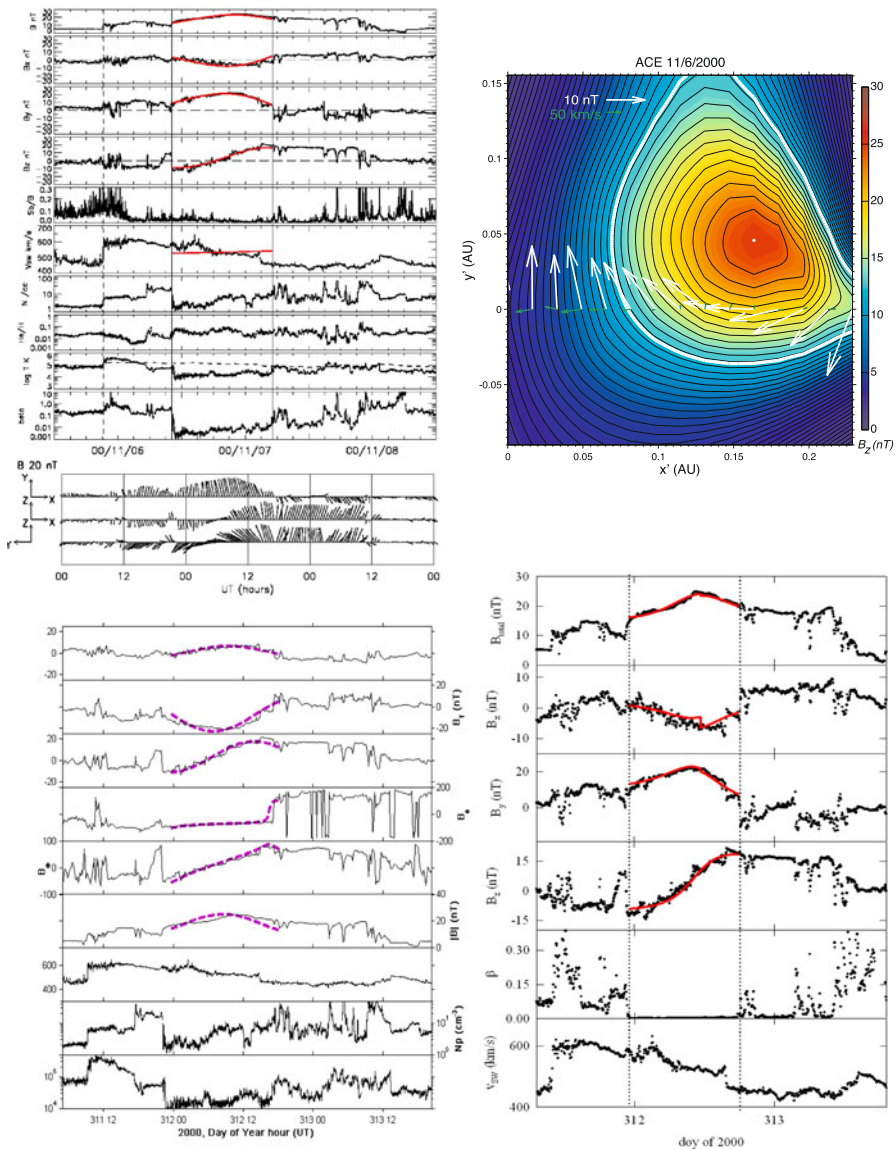


Figure 7 Same as Figure 5 but for event 27.

that the variability between methods for the same event was significantly lower than the variability between events for the same method (see their Figure 4).

For only one event, all four methods found an orientation of the ICME axis within $\pm 45^\circ$. Considering only the two force-free fitting methods (with and without expansion) and the Grad–Shafranov reconstruction method, this is the case for about 30 % of the ICMEs (seven out of the 20 common cases). Directly comparing the two force-free constant- α codes (with and without expansion), they return ICME axis within $\pm 30^\circ$ in only $\sim 30\%$ of the cases and within $\pm 45^\circ$ in only $\sim 45\%$ of the cases. While these two methods assumed very sim-

ilar geometries for the ICME, the fitting procedures are very different: a minimum-variance analysis was used to determine the axis' direction for the FFCAs code, whereas a fitting procedure was used for the expanding cylinder method. This is likely to account for the difference in the direction of the ICME axis. Finally, the elliptical cross-section model typically gives different orientation as compared to the other methods, which agrees with at least one other method for only about 30 % of the events. These results were found by letting researchers select their own boundaries.

When only two methods successfully reconstructed an event, the orientation found by these two methods usually disagreed (by more than $\pm 45^\circ$ in 15 out of the 18 cases). When the two force-free constant- α fitting methods and the Grad-Shafranov reconstruction method successfully reconstructed an event, the orientation from the Grad-Shafranov method was within $\pm 30^\circ$ of that from one of the force-free constant- α methods in 70 % of the cases (14 out of 20 events). Combining all four methods, more than 65 % (23/34) of the events reconstructed by three or more methods had at least two methods that yielded an orientation within $\pm 30^\circ$. It is therefore clear that having multiple methods able to successfully fit or reconstruct the same event gives more reliable results regarding the orientation of the ICME axis.

We furthermore quantified the importance of the boundary selection by performing the analyses again for three events using the same boundaries for all methods. This way we found a better agreement between the different codes regarding the ICME magnetic field strength, and, for one event, a better agreement for the ICME orientation.

Finally, for two methods we compared the impact parameter and magnetic field strength of MCs and non-MC ejecta. We found that the non-MC ejecta have, as expected, a weaker magnetic field but we also found some evidence that the impact parameter is, on average, larger for non-MC ejecta than for MCs. This is a first statistical hint that most (or all) ICMEs observed at 1 AU may in fact be magnetic flux ropes, but depending on how they impact the observing spacecraft, they may be recognized as magnetic clouds or not. After performing fitting and reconstructions for 59 ICMEs with four different codes, we will investigate the nature of the magnetic fields in ICMEs in more detail and also aim improve the reconstruction methods in the future.

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References

- Al-Haddad, N., Roussev, I.I., Möstl, C., Jacobs, C., Lugaz, N., Poedts, S., Farrugia, C.J.: 2011, *Astrophys. J. Lett.* **738**, L18. doi:[10.1088/2041-8205/738/2/L18](https://doi.org/10.1088/2041-8205/738/2/L18).
- Bothmer, V., Schwenn, R.: 1998, *Ann. Geophys.* **16**, 1. doi:[10.1007/s00585-997-0001-x](https://doi.org/10.1007/s00585-997-0001-x).
- Burlaga, L.F.: 1988, *J. Geophys. Res.* **93**, 7217. doi:[10.1029/JA093iA07p07217](https://doi.org/10.1029/JA093iA07p07217).
- Burlaga, L.F., Plunkett, S.P., St. Cyr, O.C.: 2002, *J. Geophys. Res.* **107**, 1266.
- Burlaga, L., Sittler, E., Mariani, F., Schwenn, R.: 1981, *J. Geophys. Res.* **86**, 6673.

- Chané, E., van der Holst, B., Jacobs, C., Poedts, S., Kimpe, D.: 2006, *Astron. Astrophys.* **447**, 727. doi:[10.1051/0004-6361:20053802](https://doi.org/10.1051/0004-6361:20053802).
- Dasso, S., Mandrini, C.H., Démoulin, P., Farrugia, C.J.: 2003, *J. Geophys. Res.* **108**, 1362. doi:[10.1029/2003JA009942](https://doi.org/10.1029/2003JA009942).
- Dasso, S., Mandrini, C.H., Démoulin, P., Luoni, M.L.: 2006, *Astron. Astrophys.* **455**, 349. doi:[10.1051/0004-6361:20064806](https://doi.org/10.1051/0004-6361:20064806).
- Dasso, S., Nakwacki, M.S., Démoulin, P., Mandrini, C.H.: 2007, *Solar Phys.* **244**, 115.
- Démoulin, P., Dasso, S.: 2009, *Astron. Astrophys.* **507**, 969. doi:[10.1051/0004-6361/200912645](https://doi.org/10.1051/0004-6361/200912645).
- Farrugia, C.J., Osherovich, V.A., Burlaga, L.F.: 1995, *J. Geophys. Res.* **100**, 12293. doi:[10.1029/95JA00272](https://doi.org/10.1029/95JA00272).
- Farrugia, C.J., Burlaga, L.F., Osherovich, V.A., Lepping, R.P.: 1992, In: Marsch, E., Schwenn, R. (eds.) *Solar Wind Seven Colloquium*, 611.
- Farrugia, C.J., Burlaga, L.F., Osherovich, V.A., Richardson, I.G., Freeman, M.P., Lepping, R.P., Lazarus, A.J.: 1993, *J. Geophys. Res.* **98**, 7621. doi:[10.1029/92JA02349](https://doi.org/10.1029/92JA02349).
- Farrugia, C.J., Berdichevsky, D.B., Möstl, C., Galvin, A.B., Leitner, M., Popecki, M.A., Simunac, K.D.C., Opitz, A., Lavraud, B., Ogilvie, K.W., Veronig, A.M., Temmer, M., Luhmann, J.G., Sauvaud, J.A.: 2011, *J. Atmos. Solar-Terr. Phys.* **73**, 1254. doi:[10.1016/j.jastp.2010.09.011](https://doi.org/10.1016/j.jastp.2010.09.011).
- Gold, T.: 1959, *J. Geophys. Res.* **64**, 1219. doi:[10.1029/JZ064i009p01219](https://doi.org/10.1029/JZ064i009p01219).
- Goldstein, H.: 1983, In: *NASA Conference Publication* **228**, 731.
- Gopalswamy, N., Akiyama, S., Yashiro, S., Michalek, G., Lepping, R.P.: 2008, *J. Atmos. Solar-Terr. Phys.* **70**, 245. doi:[10.1016/j.jastp.2007.08.070](https://doi.org/10.1016/j.jastp.2007.08.070).
- Gopalswamy, N., Xie, H., Mäkelä, P., Akiyama, S., Yashiro, S., Kaiser, M.L., Howard, R.A., Bougeret, J.L.: 2010, *Astrophys. J.* **710**, 1111. doi:[10.1088/0004-637X/710/2/1111](https://doi.org/10.1088/0004-637X/710/2/1111).
- Gosling, J.T.: 1990, In: *AGU Geophys. Monogr. Ser.* **58**, 343.
- Gulisano, A.M., Dasso, S., Mandrini, C.H., Démoulin, P.: 2005, *J. Atmos. Solar-Terr. Phys.* **67**, 1761. doi:[10.1016/j.jastp.2005.02.026](https://doi.org/10.1016/j.jastp.2005.02.026).
- Hidalgo, M.A.: 2003, *J. Geophys. Res.* **108**, 1320. doi:[10.1029/2002JA009818](https://doi.org/10.1029/2002JA009818).
- Hidalgo, M.A., Nieves-Chinchilla, T., Cid, C.: 2002, *Geophys. Res. Lett.* **29**(13), 1637. doi:[10.1029/2001GL013875](https://doi.org/10.1029/2001GL013875).
- Hu, Q., Sonnerup, B.U.Ö.: 2001, *Geophys. Res. Lett.* **28**, 467. doi:[10.1029/2000GL012232](https://doi.org/10.1029/2000GL012232).
- Hu, Q., Sonnerup, B.U.Ö.: 2002, *J. Geophys. Res.* **107**, 1142. doi:[10.1029/2001JA000293](https://doi.org/10.1029/2001JA000293).
- Hu, Q., Smith, C.W., Ness, N.F., Skoug, R.M.: 2004, *J. Geophys. Res.* **109**, A03102. doi:[10.1029/2003JA010101](https://doi.org/10.1029/2003JA010101).
- Isavnin, A., Kilpua, E.K.J., Koskinen, H.E.J.: 2011, *Solar Phys.* **273**, 205. doi:[10.1007/s11207-011-9845-z](https://doi.org/10.1007/s11207-011-9845-z).
- Jian, L., Russell, C.T., Luhmann, J.G., Skoug, R.M.: 2006, *Solar Phys.* **239**, 393. doi:[10.1007/s11207-006-0133-2](https://doi.org/10.1007/s11207-006-0133-2).
- Kahler, S.W., Krucker, S., Szabo, A.: 2011, *J. Geophys. Res.* **116**, A01104. doi:[10.1029/2010JA015328](https://doi.org/10.1029/2010JA015328).
- Kilpua, E.K.J., Liewer, P.C., Farrugia, C., Luhmann, J.G., Möstl, C., Li, Y., Liu, Y., Lynch, B.J., Russell, C.T., Vourlidas, A., Acuna, M.H., Galvin, A.B., Larson, D., Sauvaud, J.A.: 2009, *Solar Phys.* **254**, 325. doi:[10.1007/s11207-008-9300-y](https://doi.org/10.1007/s11207-008-9300-y).
- Kilpua, E.K.J., Lee, C.O., Luhmann, J.G., Li, Y.: 2011, *Ann. Geophys.* **29**, 1455. doi:[10.5194/angeo-29-1455-2011](https://doi.org/10.5194/angeo-29-1455-2011).
- Lepping, R.P., Burlaga, L.F., Jones, J.A.: 1990, *J. Geophys. Res.* **95**, 11957.
- Lepping, R.P., Acuna, M.H., Burlaga, L.F., Farrell, W.M.: 1995, *Space Sci. Rev.* **71**, 207.
- Lepping, R.P., Berdichevsky, D.B., Wu, C.C., Szabo, A., Narock, T., Mariani, F., Lazarus, A.J., Quivers, A.J.: 2006, *Ann. Geophys.* **24**, 215. doi:[10.5194/angeo-24-215-2006](https://doi.org/10.5194/angeo-24-215-2006).
- Li, Y., Luhmann, J.G., Lynch, B.J., Kilpua, E.K.J.: 2011, *Solar Phys.* **270**, 331. doi:[10.1007/s11207-011-9722-9](https://doi.org/10.1007/s11207-011-9722-9).
- Liu, Y., Luhmann, J.G., Huttunen, K.E.J., Lin, R.P., Bale, S.D., Russell, C.T., Galvin, A.B.: 2008, *Astrophys. J. Lett.* **677**, 133. doi:[10.1086/587839](https://doi.org/10.1086/587839).
- Lugaz, N., Manchester, W.B., Gombosi, T.I.: 2005a, *Astrophys. J.* **627**, 1019.
- Lugaz, N., Manchester, W.B., Gombosi, T.I.: 2005b, *Astrophys. J.* **634**, 651.
- Lugaz, N., Vourlidas, A., Roussev, I.I., Morgan, H.: 2009, *Solar Phys.* **256**, 269. doi:[10.1007/s11207-009-9339-4](https://doi.org/10.1007/s11207-009-9339-4).
- Lundquist, S.: 1950, *Ark. Fys.* **2**, 361.
- Lynch, B.J., Zurbuchen, T.H., Fisk, L.A., Antiochos, S.K.: 2003, *J. Geophys. Res.* **108**, 1239. doi:[10.1029/2002JA009591](https://doi.org/10.1029/2002JA009591).
- Manchester, W.B., Gombosi, T.I., Roussev, I., Ridley, A., De Zeeuw, D.L., Sokolov, I.V., Powell, K.G., Tóth, G.: 2004, *J. Geophys. Res.* **109**, A02107.
- Manchester, W.B., Vourlidas, A., Tóth, G., Lugaz, N., Roussev, I.I., Sokolov, I.V., Gombosi, T.I., De Zeeuw, D.L., Opher, M.: 2008, *Astrophys. J.* **684**, 1448. doi:[10.1086/590231](https://doi.org/10.1086/590231).
- Marubashi, K.: 1986, *Adv. Space Res.* **6**, 335. doi:[10.1016/0273-1177\(86\)90172-9](https://doi.org/10.1016/0273-1177(86)90172-9).

- Marubashi, K.: 2000, *Adv. Space Res.* **26**, 55. doi:[10.1016/S0273-1177\(99\)01026-1](https://doi.org/10.1016/S0273-1177(99)01026-1).
- Marubashi, K., Lepping, R.P.: 2007, *Ann. Geophys.* **25**, 2453. doi:[10.5194/angeo-25-2453-2007](https://doi.org/10.5194/angeo-25-2453-2007).
- McComas, D.J., Bame, S.J., Barker, P., Feldman, W.C.: 1998, *Space Sci. Rev.* **86**, 563.
- Möstl, C., Miklenic, C., Farrugia, C.J., Temmer, M., Veronig, A., Galvin, A.B., Vršnak, B., Biernat, H.K.: 2008, *Ann. Geophys.* **26**, 3139. doi:[10.5194/angeo-26-3139-2008](https://doi.org/10.5194/angeo-26-3139-2008).
- Möstl, C., Farrugia, C.J., Miklenic, C., Temmer, M., Galvin, A.B., Luhmann, J.G., Kilpua, E.K.J., Leitner, M., Nieves-Chinchilla, T., Veronig, A., Biernat, H.K.: 2009a, *J. Geophys. Res.* **114**, A04102. doi:[10.1029/2008JA013657](https://doi.org/10.1029/2008JA013657).
- Möstl, C., Farrugia, C.J., Biernat, H.K., Leitner, M., Kilpua, E.K.J., Galvin, A.B., Luhmann, J.G.: 2009b, *Solar Phys.* **256**, 427. doi:[10.1007/s11207-009-9360-7](https://doi.org/10.1007/s11207-009-9360-7).
- Möstl, C., Temmer, M., Rollett, T., Farrugia, C.J., Liu, Y., Veronig, A.M., Leitner, M., Galvin, A.B., Biernat, H.K.: 2010, *Geophys. Res. Lett.* **37**, L24103. doi:[10.1029/2010GL045175](https://doi.org/10.1029/2010GL045175).
- Mulligan, T., Russell, C.T.: 2001, *J. Geophys. Res.* **106**, 10581. doi:[10.1029/2000JA900170](https://doi.org/10.1029/2000JA900170).
- Nelder, J.A., Mead, R.: 1965, *Comput. J.* **7**, 308.
- Nieves-Chinchilla, T., Hidalgo, M.A., Sequeiros, J.: 2005, *Solar Phys.* **232**, 105.
- Nieves-Chinchilla, T., Colaninno, R., Vourlidis, A., Szabo, A., Lepping, R.P., Boardsen, S.A., Anderson, B.J., Korth, H.: 2012, *J. Geophys. Res.* **117**, A06106.
- Odstřcil, D., Pizzo, V.J.: 1999, *J. Geophys. Res.* **104**(13), 483.
- Odstřcil, D., Pizzo, V.J.: 2009, *Solar Phys.* **259**, 297. doi:[10.1007/s11207-009-9449-z](https://doi.org/10.1007/s11207-009-9449-z).
- Ogilvie, K.W., Chornay, D.J., Fritzenreiter, R.J., Hunsaker, F.: 1995, *Space Sci. Rev.* **71**, 55.
- Owens, M.J., Merkin, V.G., Riley, P.: 2006, *J. Geophys. Res.* **111**, A03104. doi:[10.1029/2005JA011460](https://doi.org/10.1029/2005JA011460).
- Owens, M.J., Démoulin, P., Savani, N.P., Lavraud, B., Ruffenach, A.: 2012, *Solar Phys.*, **278**, 435. doi:[10.1007/s11207-012-9939-2](https://doi.org/10.1007/s11207-012-9939-2).
- Rees, A., Forsyth, R.J.: 2004, *Geophys. Res. Lett.* **31**, L06804. doi:[10.1029/2003GL018330](https://doi.org/10.1029/2003GL018330).
- Richardson, I.G., Cane, H.V.: 2010, *Solar Phys.* **264**, 189. doi:[10.1007/s11207-010-9568-6](https://doi.org/10.1007/s11207-010-9568-6).
- Riley, P., Gosling, J.T., Pizzo, V.J.: 1997, *J. Geophys. Res.* **102**, 14677. doi:[10.1029/97JA01131](https://doi.org/10.1029/97JA01131).
- Riley, P., Linker, J.A., Mikić, Z., Odstřcil, D., Zurbuchen, T.H., Lario, D., Lepping, R.P.: 2003, *J. Geophys. Res.* **108**, 1272. doi:[10.1029/2002JA009760](https://doi.org/10.1029/2002JA009760).
- Riley, P., Linker, J.A., Lionello, R., Mikić, Z., Odstřcil, D., Hidalgo, M.A., Cid, C., Hu, Q., Lepping, R.P., Lynch, B.J., Rees, A.: 2004, *J. Atmos. Solar-Terr. Phys.* **66**, 1321. doi:[10.1016/j.jastp.2004.03.019](https://doi.org/10.1016/j.jastp.2004.03.019).
- Riley, P., Lionello, R., Mikić, Z., Linker, J.: 2008, *Astrophys. J.* **672**, 1221. doi:[10.1086/523893](https://doi.org/10.1086/523893).
- Romashets, E.P., Vandas, M.: 2003, *Geophys. Res. Lett.* **30**(20), 2065. doi:[10.1029/2003GL017692](https://doi.org/10.1029/2003GL017692).
- Rouillard, A.P.: 2011, *J. Atmos. Solar-Terr. Phys.* **73**, 1201.
- Ruffenach, A.B., Lavraud, B., Owens, M.J., et al.: 2012, *J. Geophys. Res.* doi:[10.1029/2012JA017624](https://doi.org/10.1029/2012JA017624)
- Savani, N.P., Rouillard, A.P., Davies, J.A., Owens, M.J., Forsyth, R.J., Davis, C.J., Harrison, R.A.: 2009, *Ann. Geophys.* **27**, 4349–4358. doi:[10.5194/angeo-27-4349-2009](https://doi.org/10.5194/angeo-27-4349-2009).
- Savani, N.P., Owens, M.J., Rouillard, A.P., Forsyth, R.J., Davies, J.A.: 2010, *Astrophys. J. Lett.* **714**, 128. doi:[10.1088/2041-8205/714/1/L128](https://doi.org/10.1088/2041-8205/714/1/L128).
- Shen, F., Feng, X., Wu, S.T., Xiang, C.: 2007, *J. Geophys. Res.* **112**, A06109. doi:[10.1029/2006JA012164](https://doi.org/10.1029/2006JA012164).
- Shimazu, H., Vandas, M.: 2002, *Earth Planets Space* **54**, 783.
- Smith, C.W., L'Heureux, J., Ness, N.F., Acuna, M.H.: 1998, *Space Sci. Rev.* **86**, 613.
- Sonnerup, B.U.Ö., Scheible, M.: 1998, *ISSI Sci. Rep. Ser.* **1**, 185.
- Steed, K., Owen, C.J., Démoulin, P., Dasso, S.: 2011, *J. Geophys. Res.* **116**, A01106. doi:[10.1029/2010JA015940](https://doi.org/10.1029/2010JA015940).
- Vandas, M., Romashets, E.P.: 2003, *Astron. Astrophys.* **398**, 801. doi:[10.1051/0004-6361:20021691](https://doi.org/10.1051/0004-6361:20021691).
- Vandas, M., Romashets, E.P., Watari, S., Geranios, A., Antoniadou, E., Zacharopoulou, O.: 2006, *Adv. Space Res.* **38**, 441. doi:[10.1016/j.asr.2004.11.026](https://doi.org/10.1016/j.asr.2004.11.026).
- Wang, Y.M., Ye, P.Z., Wang, S.: 2003, *J. Geophys. Res.* **108**, 1370.
- Yurchyshyn, V.B., Wang, H., Goode, P.R., Deng, Y.: 2001, *Astrophys. J.* **563**, 381. doi:[10.1086/323778](https://doi.org/10.1086/323778).
- Zhang, J., Richardson, I.G., Webb, D.F., Gopalswamy, N., Huttunen, E., Kasper, J.C., Nitta, N.V., Poomvises, W., Thompson, B.J., Wu, C.C., Yashiro, S., Zhukov, A.N.: 2007, *J. Geophys. Res.* **112**, A10102. doi:[10.1029/2007JA012321](https://doi.org/10.1029/2007JA012321).
- Zhao, X.P., Hoeksema, J.T.: 1998, *J. Geophys. Res.* **103**, 2077. doi:[10.1029/97JA03234](https://doi.org/10.1029/97JA03234).
- Zurbuchen, T.H., Richardson, I.G.: 2006, *Space Sci. Rev.* **123**, 31. doi:[10.1007/s11214-006-9010-4](https://doi.org/10.1007/s11214-006-9010-4).